April 1932

## POULKOVO OBSERVATORY CIRCULAR

## **№** 2

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## Note on the Continuous Spectrum of Solar Faculae.

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In spring 1931 we obtained with the 30 refractor a series of solar faculae pectrograms by means of a three-prism spectrograph. The plates were standartized by means of a light-reducer. However, only on one plate (taken on May 19, 1931) the image of the facula observed in the western part of the solar disk at a distance of 0.94 R from the centre was sufficiently sharp. On the remaining plates the images of the faculae were hardly discernible in the solar spectrum forming their background

The intensities of the continuous faculae spectrum were measured in four places by comparing them to the intensities of the non-perturbed regions of the solar spectrum on both sides of the facula along the spectrograph slit. Whence by interpolation the intensity of the facula spectrum was determined relatively to the spectrum of the solar disk region at the same distance from the disk centre. The following results were obtained

	λ	I
have for the photograph of	389.7	114
•	391.8	115
	498.1-	110
	506.3	110.

For each wave length the intensity of the solar spectrum is taken as 100. The above table shows a certain excess of radiation in the violet region, denoting a greater effective temperature of the facula as compared to the effective temperature of the sun's surface. The effective temperature of the facula is only by 100° higher than that of the solar surface at the same distance from the centre.

The faculae apparently represent such regions of the sun's surface where a change from a state of monochromatic radiative equilibrium to a state of local thermodynamic equilibrium takes place in higher layers. Only that circumstance enables us to explain the fact that faculae are more distinctly seen at the edge of the solar disk, and that the number of atoms obtained from  $Ca^+$  absorption lines are less than the corresponding numbers obtained from the normal solar spectrum. However the question why the faculae give a greater energy flux than the sun's surface, requires a special investigation from the theoretical standpoint.

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